All CCDs on HST are subject to radiation damage, which deteriorates the detector performance. One of the most severe effects is the degradation of the charge transfer efficiency (CTE). Poor CTE has the effect of removing small amount of charge from the charge packets during the transfer from one pixel to the next severely affecting the accuracy of the photometry.

Correction to restore measured integrated counts to their "true" value is therefore required.

The charge loss due to imperfect CTE depends on the combination of many parameters, such as:
- the position of the source on the detector,
- the total counts of the source,
- the physical extent of the source,
- the level of the background,
- the amount of damages suffered by the detector.

Riess and Mack 2004 provided the following CTE correction for WFC data:

\[
Y_{CTE} = 10^{A \times \text{SKY} + B \times \text{FLUX} + C \times \frac{Y}{2048} \times (MJD - 52333) / 365}
\]

where coefficients A, B, and C depend on the aperture size adopted for the photometry and Y represents the number of transfer in the parallel direction.

The proposed Y_{CTE} parameterization can not be directly applied to complex mosaicked drizzled data, as the one shown in Figure 1, where any given star can be at very different distances from the amplifiers, and thus be subject to different amount of charge losses in each of the individual frames.

We solved this problem applying the following steps:
- Perform the photometric analysis of the entire mosaicked image; this provides the most accurate magnitude (\text{mag}) determination for the stars detected in the image.
- Correct each individual frame for geometric distortion and perform core aperture photometry on the detected stars.
- Calculate the Y_{CTE} correction for the photometry obtained on the single exposure.
- Cross-correlate all the catalogs obtained from the aperture photometry and compute for each detected source the average value of Y_{CTE}. The single exposures are still affected by cosmic rays; to obtain a clean determination of \langle Y_{CTE} \rangle a sigma clipping on the associate magnitudes may be required.
- The average correction is then applied to the magnitude of each star: \text{mag} + \langle Y_{CTE} \rangle.
- Given that the uncertainties on Y_{CTE} for very faint stars is comparable to the photometric errors, stars too faint to be detected on single exposures are not corrected for CTE (see Fig 2).